

Is the Fuzziness of GRB970228 constant?

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Abstract. In view of the data gathered in September 1997, we review the flux values collected so far for the "fuzziness" seen in the optical counterpart of GRB970228. Comparison between the ground based data collected in March and the data of September 1997 suggests a fading of the fuzz. Given the diversity of the data in hand, the magnitude of the effect and its significance are not easy to quantify. Only new images, both from the ground and with the Space Telescope, directly comparable to the old ones could settle this problem.

INTRODUCTION

After the SAX positioning of GRB970228 (Costa et al. 1997), and the discovery of an optical transient in the refined error box (van Paradijs et al, 1997), the optical counterpart of GRB970228 has been observed many times both with ground based instruments and with the Hubble Space Telescope. Several days after the event, an extended optical emission was detected where the Optical Transient (OT) had been seen in the discovery image, taken 21 h after the event (van Paradijs et al, 1997). Since then, the magnitude of such an extended emission has been measured many times, by several observers, using different instrumental set-ups. In this paper we review and compare the measurements gathered so far to investigate if the flux values recently measured by STIS on HST (Fruchter et al, 1997) and by the 5m Palomar telescope (Djorgovski et al., 1997) are consistent with the ground based ones obtained at early epochs.

THE DATA

Table 1 summarizes the data collected so far, both for the OT integrated magnitude (ground measurements) and for the contribution of the two components : point source and extended emission (HST data). The first claim for an extended object, using an 1 hour NTT exposure taken on March 13th, gave $m_R = 23.8 \pm 0.2$ (van Paradijs et al 1997). To this, one should add the Keck

measurement ($m_R = 24.0 \pm 0.2$, Metzger et al, 1997a) obtained on March 6th but announced after the discovery of the optical transient (Groot et al. 1997a).

Indeed, all the March ground measurements agree in describing the extended object as elongated in the North-South direction with $m_R \sim 24$. HST observations were carried out in late March and early April using the WFPC and a broad V filter (F606). The extended source was resolved into a point source superimposed to a "fuzz", which, according to Sahu et al. (1997), were detected at $m_V = 26.1 \pm 0.1$ and $m_V = 24.9 \pm 0.3$, respectively. Comparison of the March and April images showed that the point source was most probably fading, while nothing definite could be said on the diffuse emission (Sahu et al. 1997). Using the same data, Galama et al (1997) estimate an R magnitude of $m_R = 25.17 \pm 0.13$ for the point source and $m_R = 24.7 \pm 0.30$ for the fuzz. In the same paper, the value of the magnitude measured by the NTT on March 13th was also revised, bringing it to $m_R = 24.3 \pm 0.2$. More Keck observations taken on April 5th and 6th gave, for the total emission, a $m_R = 24.9 \pm 0.3$ (Metzger et al. 1997b) i.e. significantly lower than both the HST one and the Keck March 6th data. In April the source became unobservable from the ground and from HST. The observability window opened again in late August when it was pointed both from the Keck/Palomar (IAU Circ 6732) and from HST using, this time, the newly installed STIS. Both observations show the overall flux to be lower than that measured previously. On Sept 4th, Djorgovski et al (1997) used the Palomar 5m telescope to obtain an R image of the field where the extended source was detected at $m_R \sim 25.5$. The STIS instrument on board HST also observed the source on Sept 4th with the Clear filter. Fruchter et al, (1997) are barely able to detect the point source, now at $m_V = 28.0 \pm 0.25$, over a diffuse emission of $m_V = 25.7 \pm 0.15$, i.e. 0.8 magnitude fainter than in the

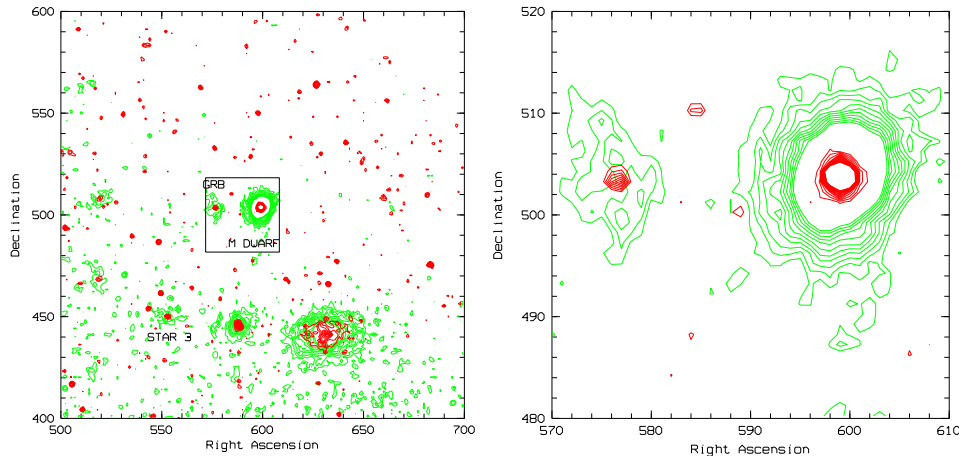


FIGURE 1. Superposition of the March HST/PC frame onto the NTT/SUSI one (actually, only a 26×26 arcsec area is shown). North to the top and East to the left. Axis units are SUSI pixels (0.13 arcsec). The GRB counterpart as well as the nearby M dwarf and star #3 (see text) are labelled. A zoom of the central square is shown in the right panel. The GRB position measured by the HST/PC falls exactly at the center of the nebulosity observed by the NTT.

<i>Date</i>	<i>Telescope</i>	<i>Filter</i>	<i>mag</i>			<i>flux</i> ($10^{-30} \text{ erg cm}^{-2} \text{ s}^{-1} \text{ Hz}^{-1}$)			<i>Ref.</i>
			<i>total</i>	<i>point source</i>	<i>extended</i>	<i>total</i>	<i>point source</i>	<i>extended</i>	
March 6 th	Keck	R	$m_R = 24 \pm 0.2$			$7.2^{+1.4}_{-1.2}$			1.
March 13 th	NTT	R	$m_R = 23.8 \pm 0.2$			$8.6^{+1.4}_{-1.7}$			2.
			$m_R = 24.3 \pm 0.2$			$5.4^{+0.9}_{-1.1}$			3.
March 26 th	HST/PC	606W	$m_R = 24.1 \pm 0.2$	$m_R = 25.17 \pm 0.13$	$m_R = 24.7 \pm 0.3$	$6.5^{+1.1}_{-1.3}$	$2.4^{+0.3}_{-0.3}$	$3.8^{+0.9}_{-1.2}$	3.
				$m_V = 26.1 \pm 0.1$	$m_V = 24.9 \pm 0.3$		$1.3^{+0.1}_{-0.1}$	$4.03^{+0.9}_{-1.3}$	4.
					$m_V = 25.6 \pm 0.25$			$2.11^{+0.4}_{-0.5}$	5.
April 6 th	Keck	R	$m_R = 24.9 \pm 0.3$			$3.11^{+0.7}_{-0.9}$			6.
April 7 th	HST/PC	606W	$m_R = 24.2 \pm 0.15$	$m_R = 25.5 \pm 0.13$	$m_R = 24.7 \pm 0.3$	$5.94^{+0.7}_{-0.8}$	$1.79^{+0.2}_{-0.2}$	$3.74^{+0.9}_{-1.2}$	3.
				$m_V = 26.4 \pm 0.1$	$m_V = 25.2 \pm 0.35$		$1.01^{+0.08}_{-0.09}$	$3.05^{+0.8}_{-1.1}$	4.
					$m_V = 25.6 \pm 0.25$			$2.11^{+0.4}_{-0.5}$	5.
Sept 4 th	5 m Hale	R	$m_R \sim 25.5$			~ 1.8			7.
Sept 4 th	HST/STIS	clear		$m_V = 28 \pm 0.25$	$m_V = 25.7 \pm 0.15$		$0.23^{+0.05}_{-0.06}$	$1.92^{+0.2}_{-0.3}$	5.

1. Metzger et al, 1997a; 2. Groot et al, 1997b (see also, van Paradijs et al, 1997); 3. Galama et al, 1997; 4. Sahu et al, 1997; 5. Fruchter et al, 1997; 6. Metzger et al, 1997b; 7. Djorgovski et al, 1997;

TABLE 1. Ground-based and HST observations of the GRB970228 optical counterpart. For each observation the telescope as well as the "original" filter are indicated. Columns 4-6 list the total magnitude of the GRB counterpart as well as the magnitudes of the point and extended source (when resolved). Corresponding monochromatic fluxes are given in column 7-9. We note that the V/R magnitudes obtained from HST broad band observations are based on ad hoc color transformation. Therefore, they are not directly comparable to the ground based R measurements which provide the long term coverage of the source evolution.

HST March observation. This prompted a reanalysis of the March/April WFPC data which resulted in a reassessment of their magnitude value now estimated at $m_V = 25.6 \pm 0.25$, i.e. half of the flux published by Sahu et al (1997) for two independent WFPC observations. Although not stated, a similar downward revision should apply also to the R magnitude values published by Galama et al (1997) for the HST observations.

However, even accepting that the HST data, after re-analysis, can be rendered consistent, it seems very difficult to reconcile the September STIS/Palomar data with the NTT/Keck ones of early March. Table 1 shows that a suggestion for fading of the extended component of the OT seems to be present. However, magnitude values do not always render easy the comparison of data taken through different filters. The suggestion for a significant fading of the extended source, implicit when comparing March to September data, becomes stronger when one computes the actual energy fluxes. This is done in Table 1, where we have transformed the magnitude values in $erg/cm^2 sec Hz$.

HST VS NTT

In the following, we shall compare the NTT data (kindly provided to us by Jan van Paradijs) with the HST ones. In order to do so we have to assume that the extended emission seen from the ground is indeed the superposition of the fuzziness seen by HST plus a point source. If we assume that the September STIS/Palomar flux values for the extended source are correct, and if we further assume no fading, we have to explain the extended total emission seen both by NTT and by Keck with a combination of the STIS/Palomar fluxes plus a point source of suitable magnitude. Even considering the revised NTT mag value given by Galama et al (1997), we have to account for a total flux of $\sim 5.4 \cdot 10^{-30} erg/cm^2 sec Hz$. Since the extended source observed in September provides $\sim 1.9 \cdot 10^{-30} erg/cm^2 sec Hz$, the unseen point source should have been $\sim 3.5 \cdot 10^{-30} erg/cm^2 sec Hz$, i.e. definitely brighter than the extended one. However, in order to simulate the appearance of such a combination one should be able to locate the HST point source into the NTT nebosity. This calls for an accurate superposition of the HST March data onto the NTT/SUSI frame. To take care of geometric distortion, we have used the task "mosaic" which re-scales the HST/PC image, rebinning it to pixel size of $0.1''$. The resulting image has been rebinned and rotated onto the SUSI one ($0.13''/\text{pixel}$) using a standard technique which is certainly accurate to better than $1/2$ pixel (actually $1/10$ would be a more realistic estimate). Figure 1 shows the superposition of the HST March frame onto the SUSI one. Only the central portion of the actual images is given. Zooming on the OT, one sees clearly that the HST point source falls in the central part of the NTT nebosity, where the emission is less intense and no hint of a point-like object is seen. However, the central region of the nebosity is just where one should put a hypothetical point source of $3.5 \cdot 10^{-30} erg/cm^2 sec Hz$, corresponding to $m_R \sim 24.65$. This value is

similar to the flux measured by the NTT for the faint source just south of the GRB970228 counterpart. Such a source is in interesting test case, since it is point-like in the HST/PC image (star #3 in Figure 1) but it looks extended in the NTT image. However, inspection of Figure 2, where we have compared the Right Ascension and Declination tracings of the two sources, shows unambiguously their difference in shape for a comparable flux. While star #3 is dominated by a clear peak superimposed to a region of higher background, the GRB nebosity does not show any obvious point-like contribution. This is somewhat surprising, since a point source of $m_R \sim 24.6$ should have been far easier to detect than a $m_R \sim 25.5$ extended one. Moreover, we note that such a faint extended source would be hardly within reach of an 1 hour NTT exposure.

Thus, the truly extended nature of the NTT source, coupled with the lack of point source at the HST location, leads to the conclusion that the nebosity itself has faded away from March 13th and Sept 4th.

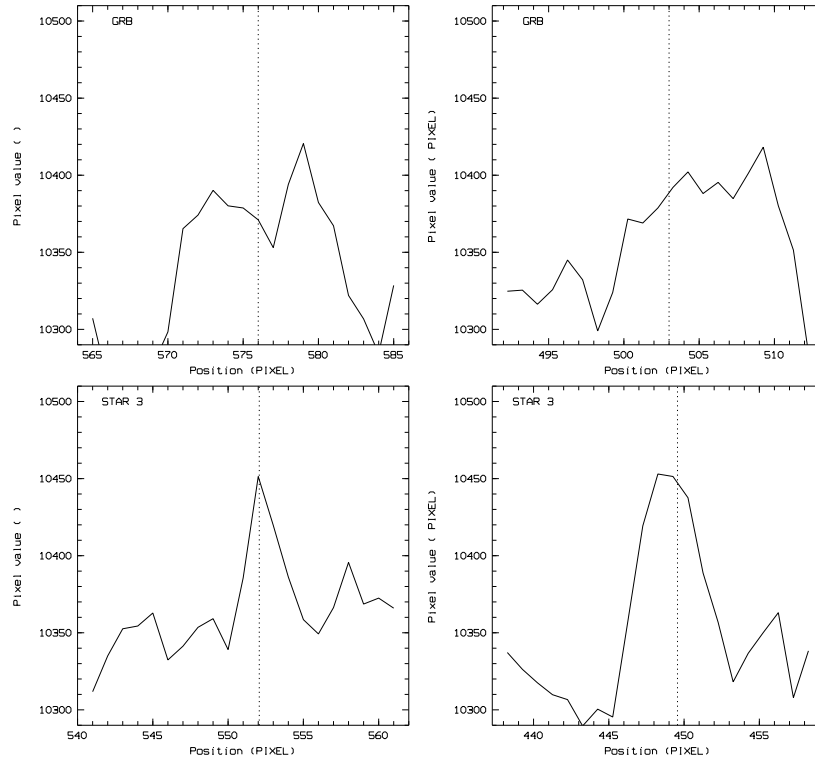


FIGURE 2. Right ascension (left) and declination (right) tracings for the GRB counterpart (upper panels) and for star #3 (lower panels) obtained from the NTT image shown in Fig.1. In both cases, the tracings have been computed at the expected location of the point sources. Their coordinates, obtained by registering the HST/PC frame onto the NTT/SUSI one, are marked as vertical lines. Although the PSF of star #3 is affected by non-optimal seeing conditions (≥ 1 arcsec), which result in a profile broadening along declination, the presence of a point source is certainly recognizable.

CONCLUSIONS

Although comparing fluxes obtained with different instruments, different filters and different observing conditions is not straightforward, the compilation of the magnitudes values measured so far for the optical counterpart of GRB970228 points toward a fading both of the point source and of the diffuse emission. While the fading of the point source is expected in all theoretical scenarios, the fading of the diffuse emission has far reaching consequences and, as such, is in need of a dedicated observing campaign. The data available are numerous, but too diverse to provide the constraints needed to assess with certainty if, and how much, the nebulosity has faded. Indeed, for GRB 970228, it looks as if every new observation results in a downward revision of the values previously published. Only more observations, directly comparable with those already in hand (i.e. obtained with identical instrumental-ups) can provide a definite answer to this all important point. Of particular importance could be new HST/PC data since the unfiltered STIS image is not directly comparable to the PC ones, obtained with a broad V filter.

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